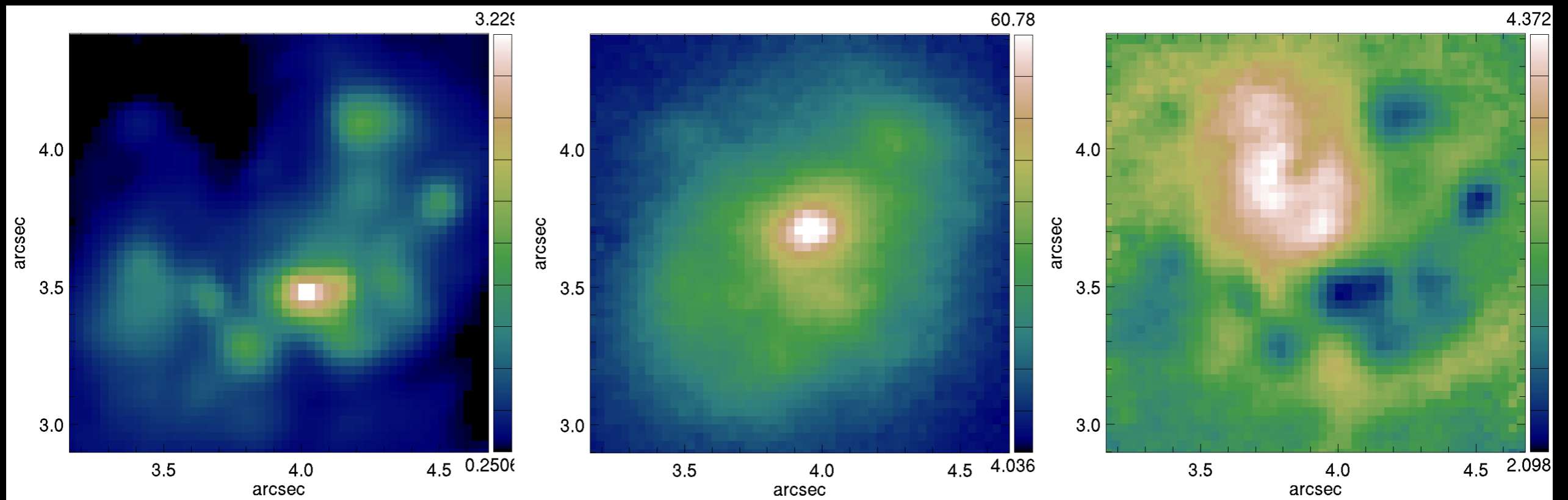


First demo science with MOAO observations of distant merging galaxies with CANARY



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(1) LESIA, Observatoire de Paris

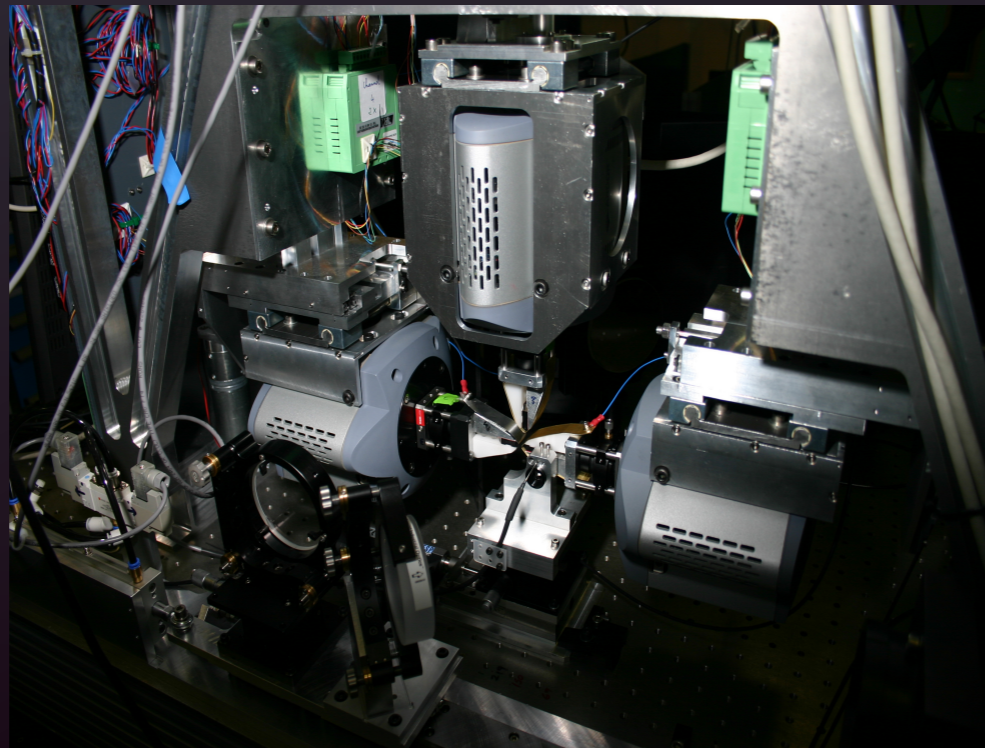
(2) University of Durham

- ◉ The CANARY demonstrator
- ◉ CAMICAz : the near-IR camera for CANARY
- ◉ The science mode of CANARY
- ◉ Observing distant galaxies
- ◉ Analysis of MOAO performance for extragalactic science
- ◉ Preliminary science results
- ◉ Conclusion and perspectives

The CANARY demonstrator

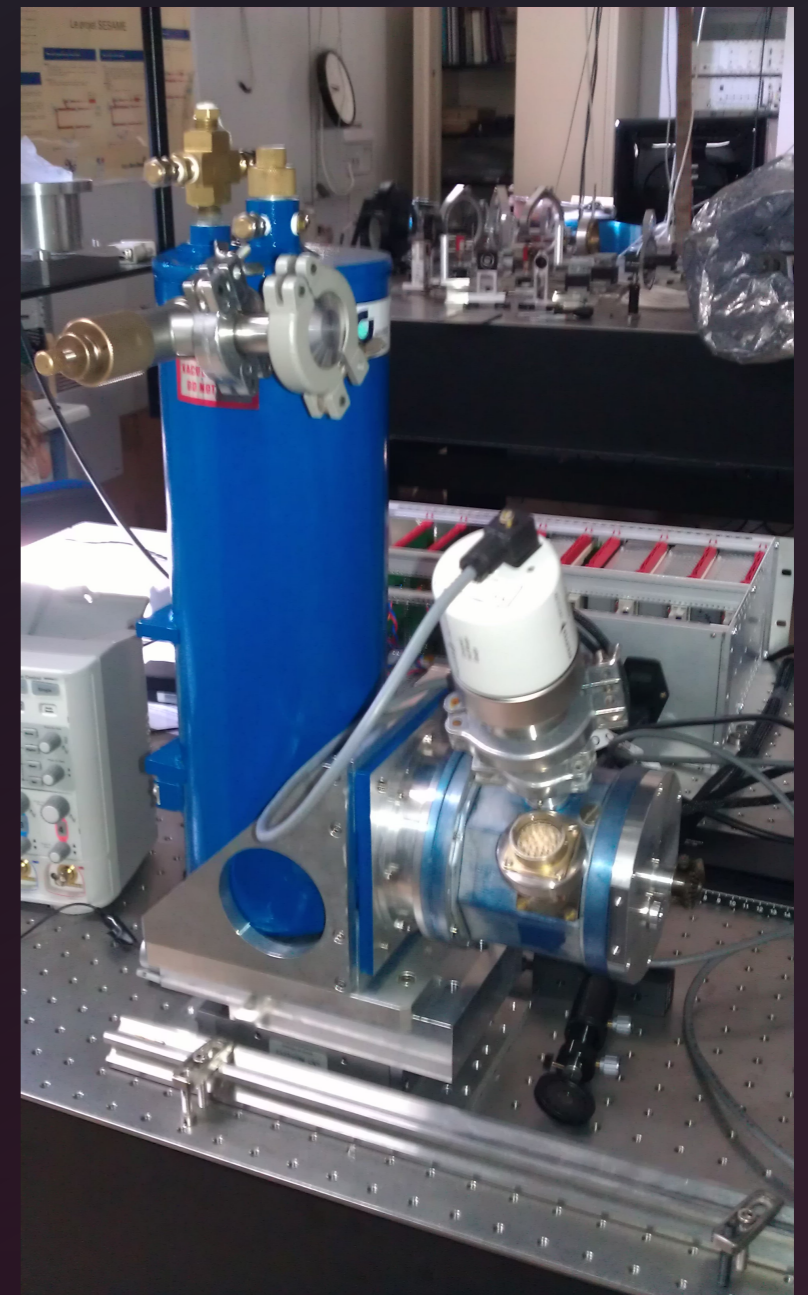


- On-sky MOAO demonstrator @WHT
 - Collaboration between University of Durham (UK) and LESIA
 - 2010-2011, Phase A : NGS tomography

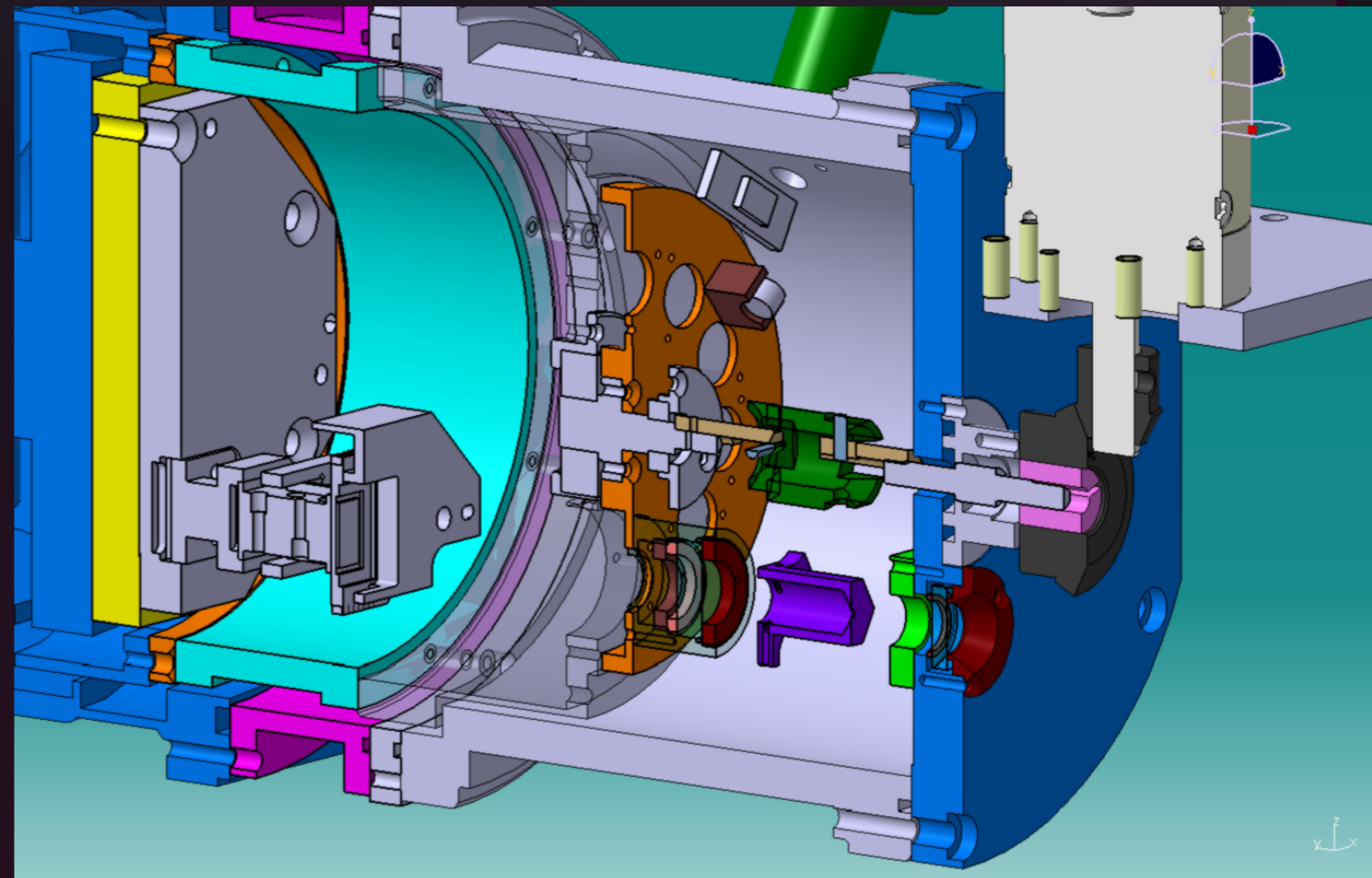


2012-2013, Phase B : mixed LGS + NGS tomography and new near-IR camera : CAMICAz

- 2014-2015, Phase C : ELT configuration and LTAO (first on-sky run in few weeks !)



- Near-IR NICMOS-based camera with dedicated cryo-cooled opto-mechanics
 - 256x256 NICMOS detector (acquired in 1991 for the now decommissioned DENIS survey)
 - Cold stop, pupil imaging lens and filter wheel
 - High performance watercooled readout electronics ($< 30e^-$ read noise)
 - 30mas pixels, J, H and K broad band filters
 - **2014 : filter wheel upgraded to 10 filter positions including narrow band filters. Funding from Région Île-de-France**



- Near-IR NICMOS-based camera with precision mechanics
 - 256x256 NICMOS detector (acquired for sky survey)
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Science mode of CANARY



- CANARY : technology demonstrator
 - Not designed to do science observations : CANARY is not a real science facility @ WHT
 - On axis beam shared by near-IR camera and truth sensor (TS) : in science mode on faint targets, TS is useless
 - Near-IR camera FoV diameter : 8"
 - Target acquisition using low sensitivity visible camera (being upgraded) : pointing can be difficult on very faint target
 - Science FoV on-sky P.A. not known precisely (lacking some interaction with the telescope infrastructure)
 - Observing block :
 - Small jitters : no need to interrupt the MOAO loop, modify reference slopes
 - Large offsets : need to interrupt the loop and offset the telescope manually

- Main concern : pseudo-static aberrations
 - Static compensation based on a combination of offsets on the DM and reference slopes on the off-axis WFS
 - Cannot be measured directly in MOAO systems without TS
 - We propose a two steps approach : determine the static component on the TS for the observing configuration using the LGS WFS measured static and the NGS in two configurations



Science mode of CANARY



- Main concern : pseudo-static aberrations
 - Observing configuration : 4 LGS WFS + 1 NGS WFS (tip-tilt +)
 - Measure static term on LGS => LGS1 and on the NGS : NGS1



NGS WFS



No TS

Science mode of CANARY



- Main concern : pseudo-static aberrations
 - Additional configuration : measure LGS2 and offsets to the DM from the TS measurements (TS2)



NGS WFS



Science mode of CANARY



- ◉ Main concern : pseudo-static aberrations
 - ◉ Back to observing conditions
 - ◉ We determine the reference slopes for the NGS WFS as :
$$\text{NGS2} = \text{NGS1} + (\text{LGS2} - \text{LGS1})$$
 - ◉ Apply offsets on the DM from TS2
 - ◉ Start observing sequence



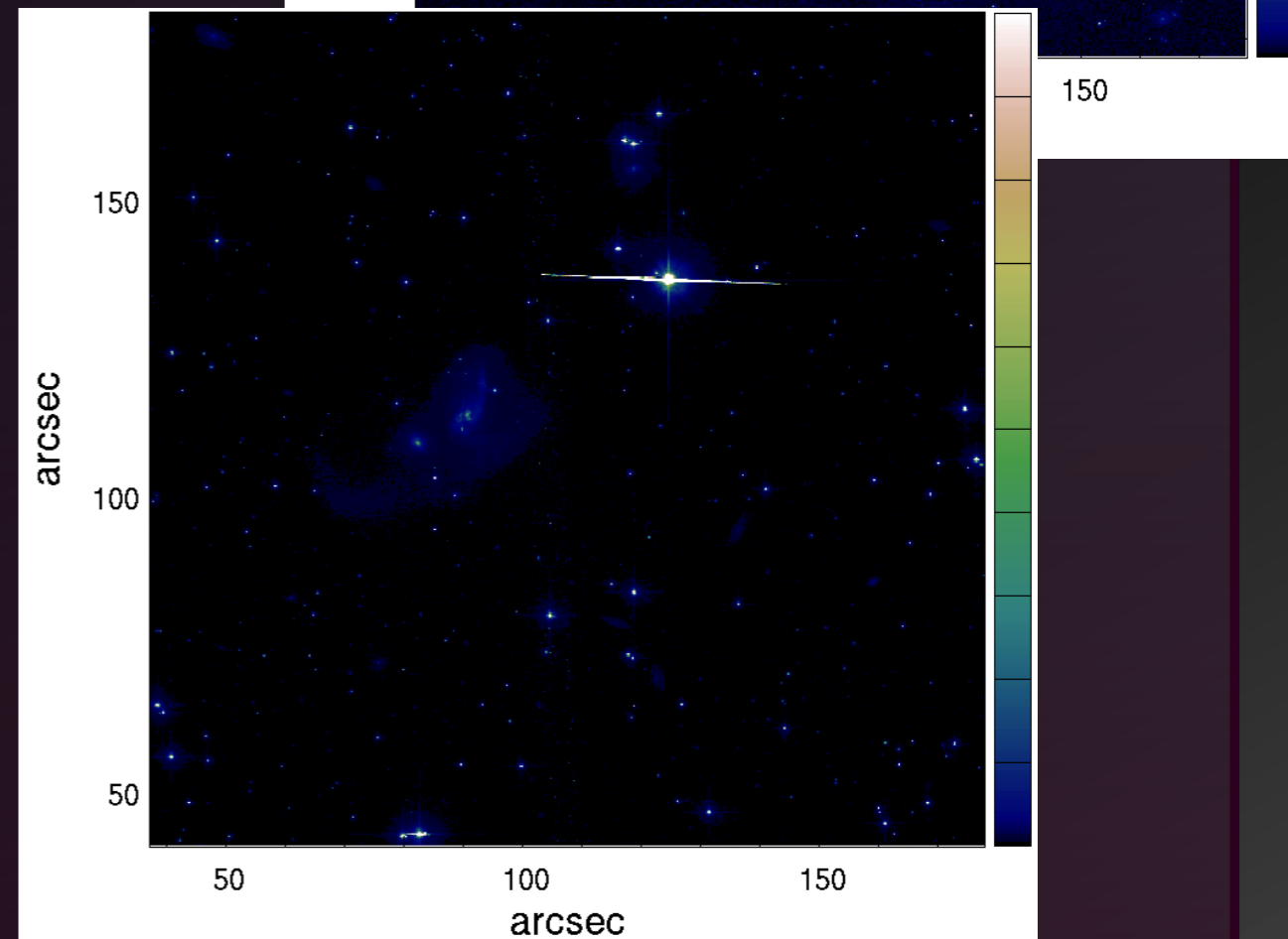
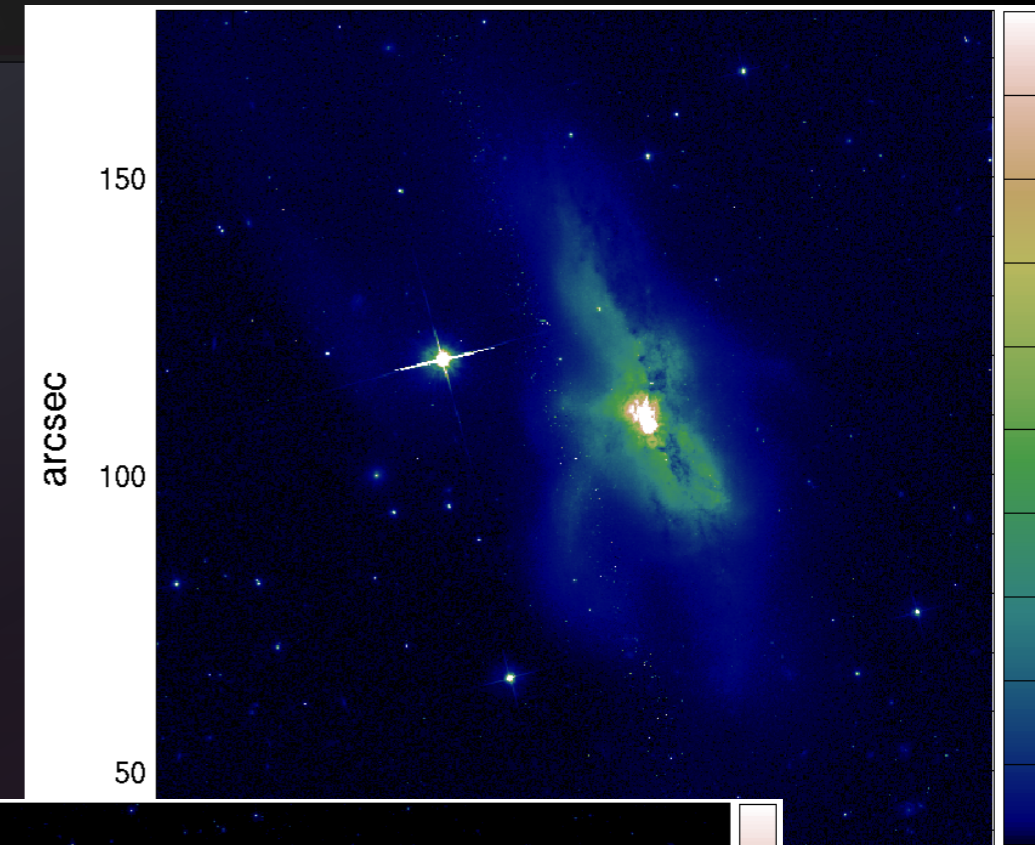
NGS WFS



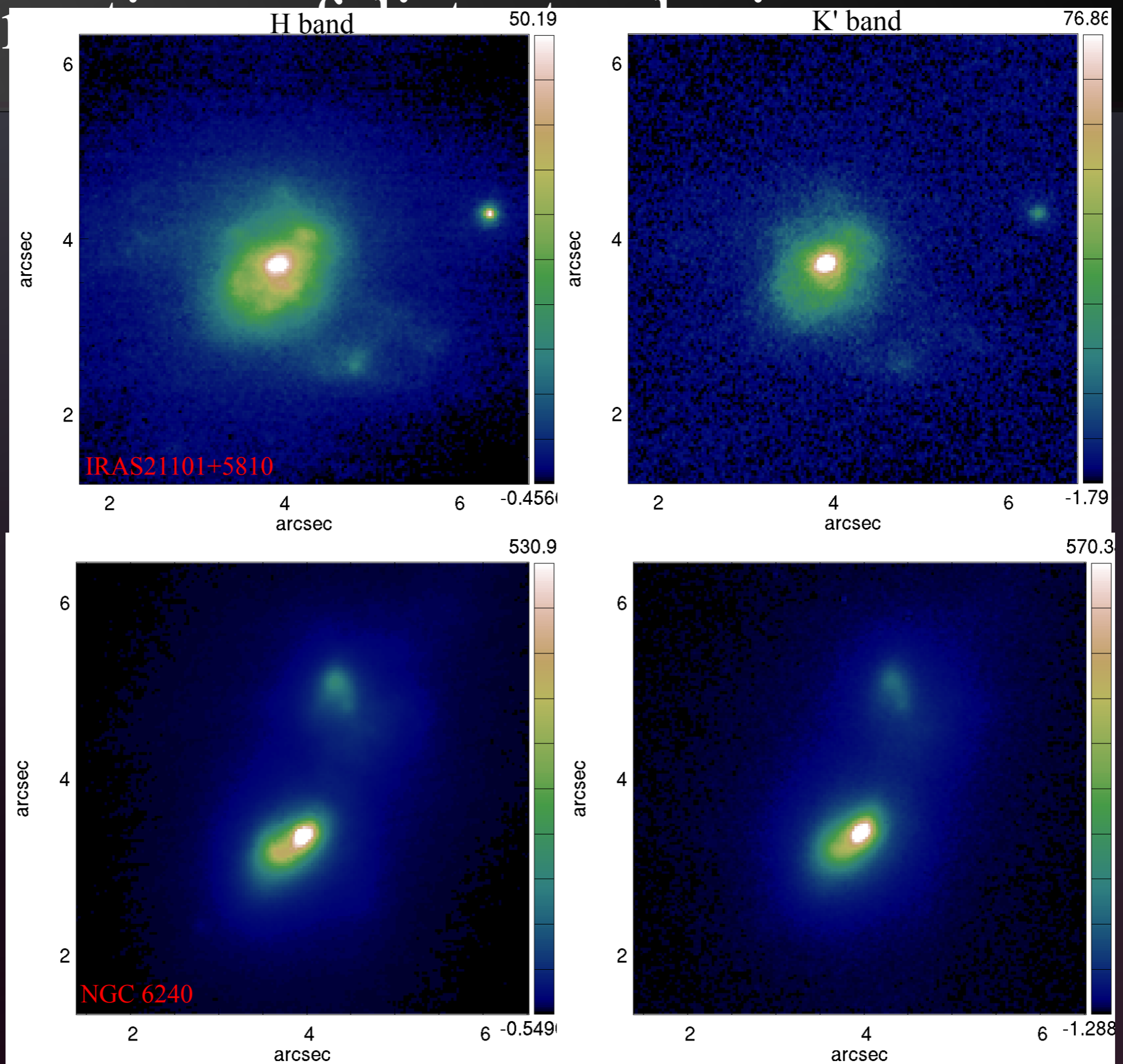
No TS

Observations of distant galaxies

- ◉ 2 merging systems from thez GOALS survey were observed in July 2013 @ broad H and K'
- ◉ NGC 6240 : $z=0.0245$
Indiv. exposures of 20s
total : 2000s per band
NGS @ 35" off-axis
- ◉ IRAS21101+5810 : $z=0.039$
Indiv. exposures of 40s
total > 1h per band
NGS @ 50" off-axis

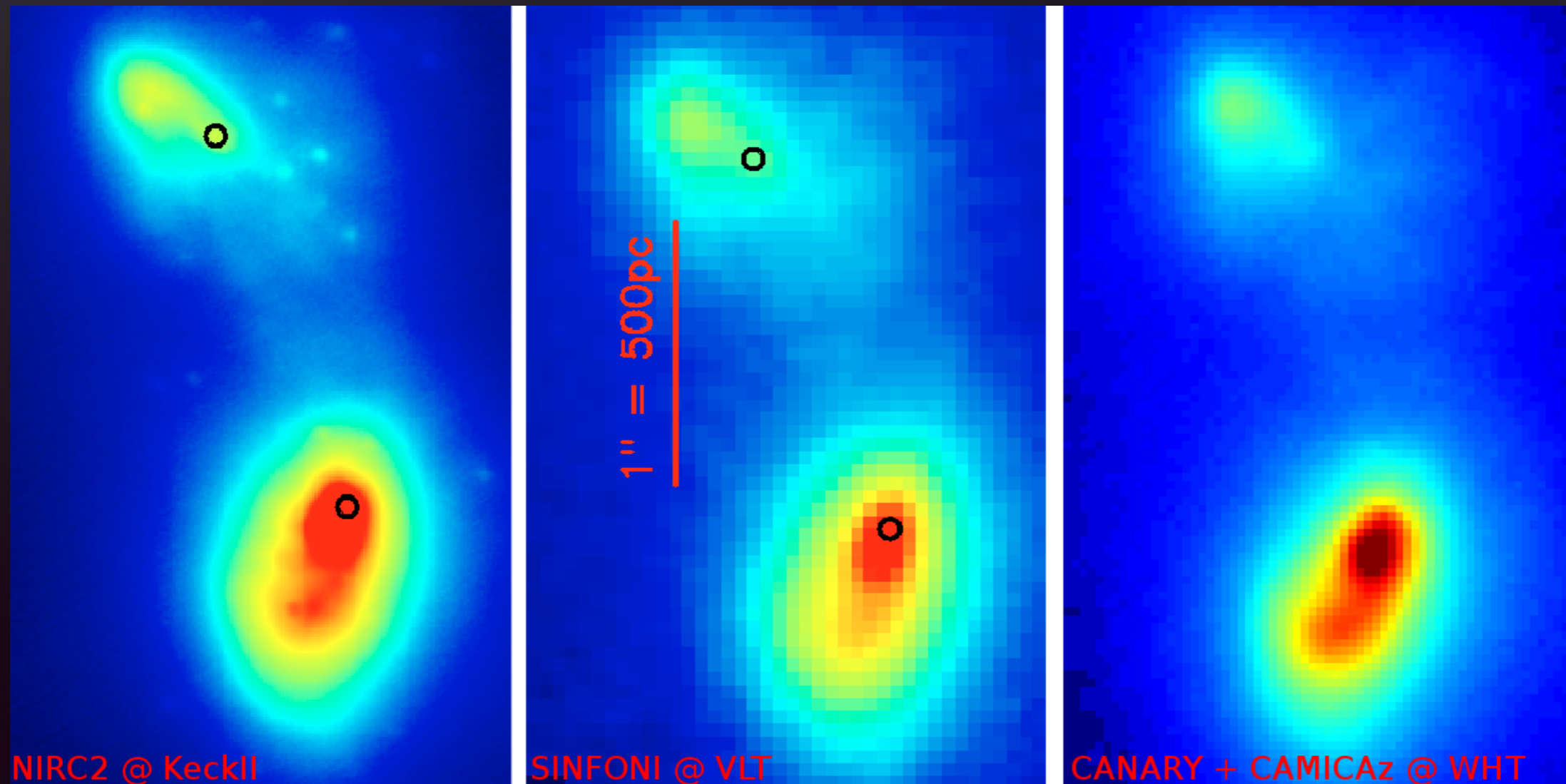


Observations



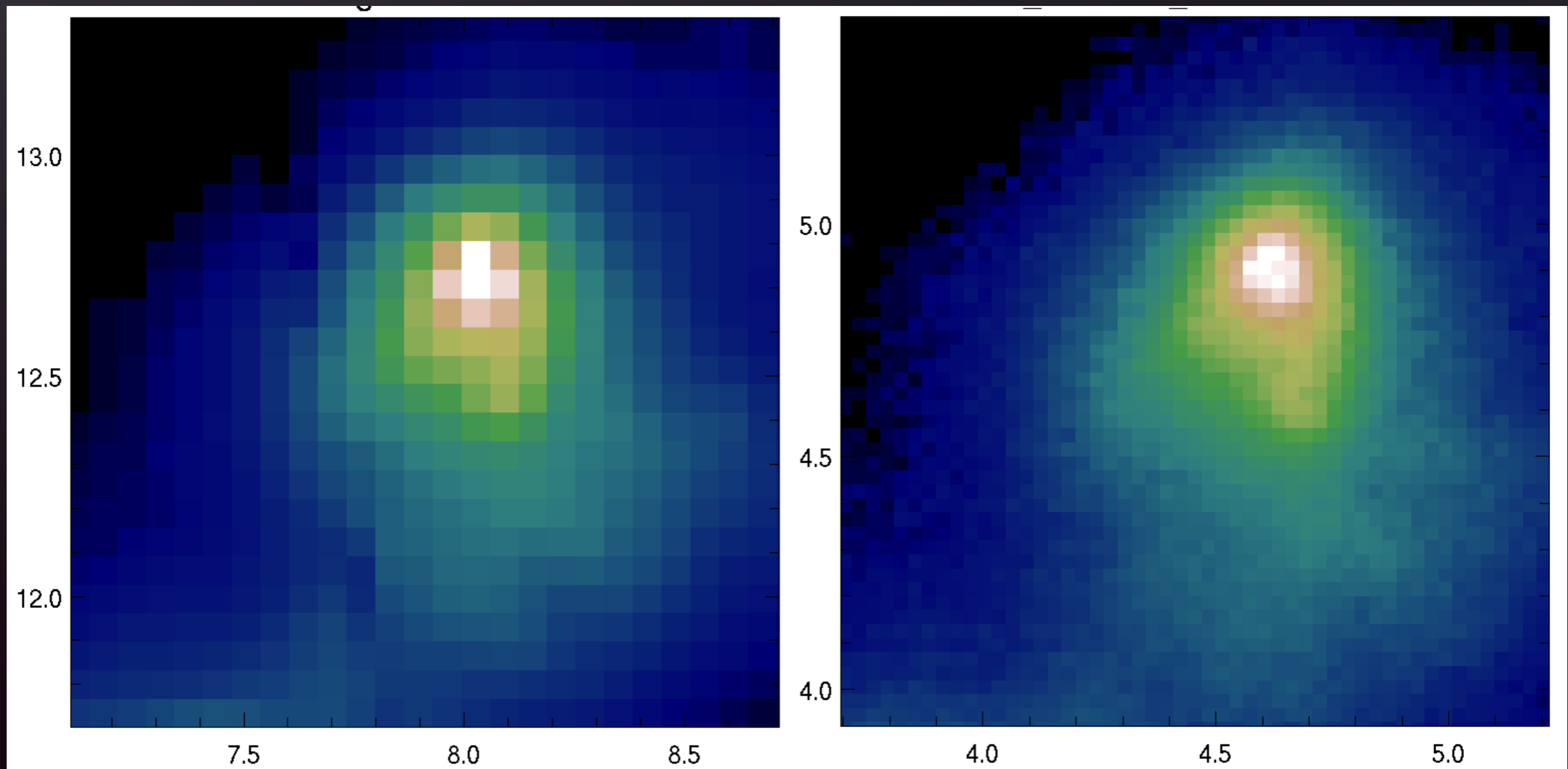
MOAO performance

- Comparing CANARY data to archive data
- NGC 6240 : already observed with LGS AO (e.g. Pollack et al. with KeckII LGS AO or Engel et al. with VLT SINFONI)



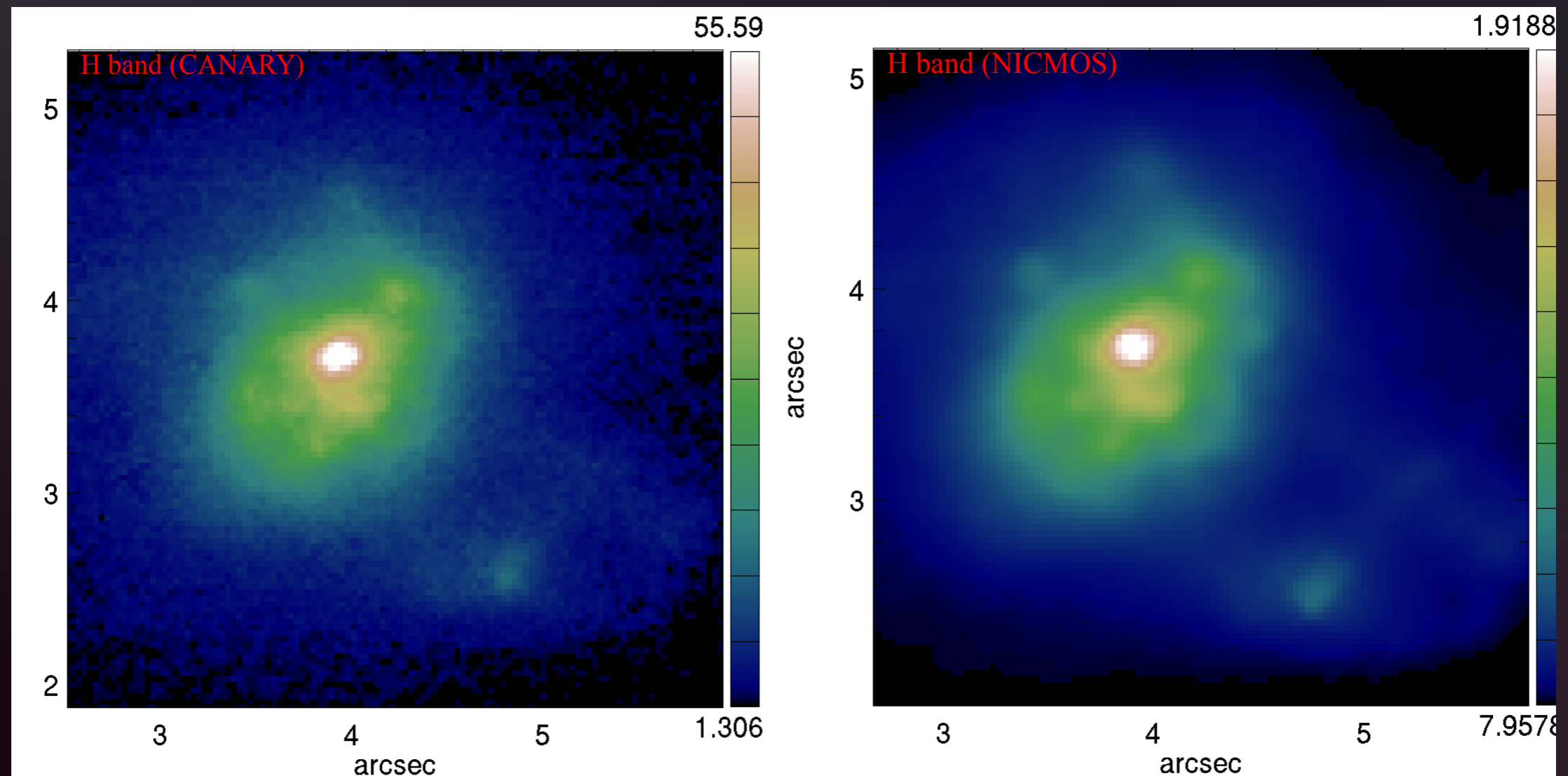
MOAO performance

- Comparing CANARY data to archive data
- NGC 6240 @ H : data from HST archive



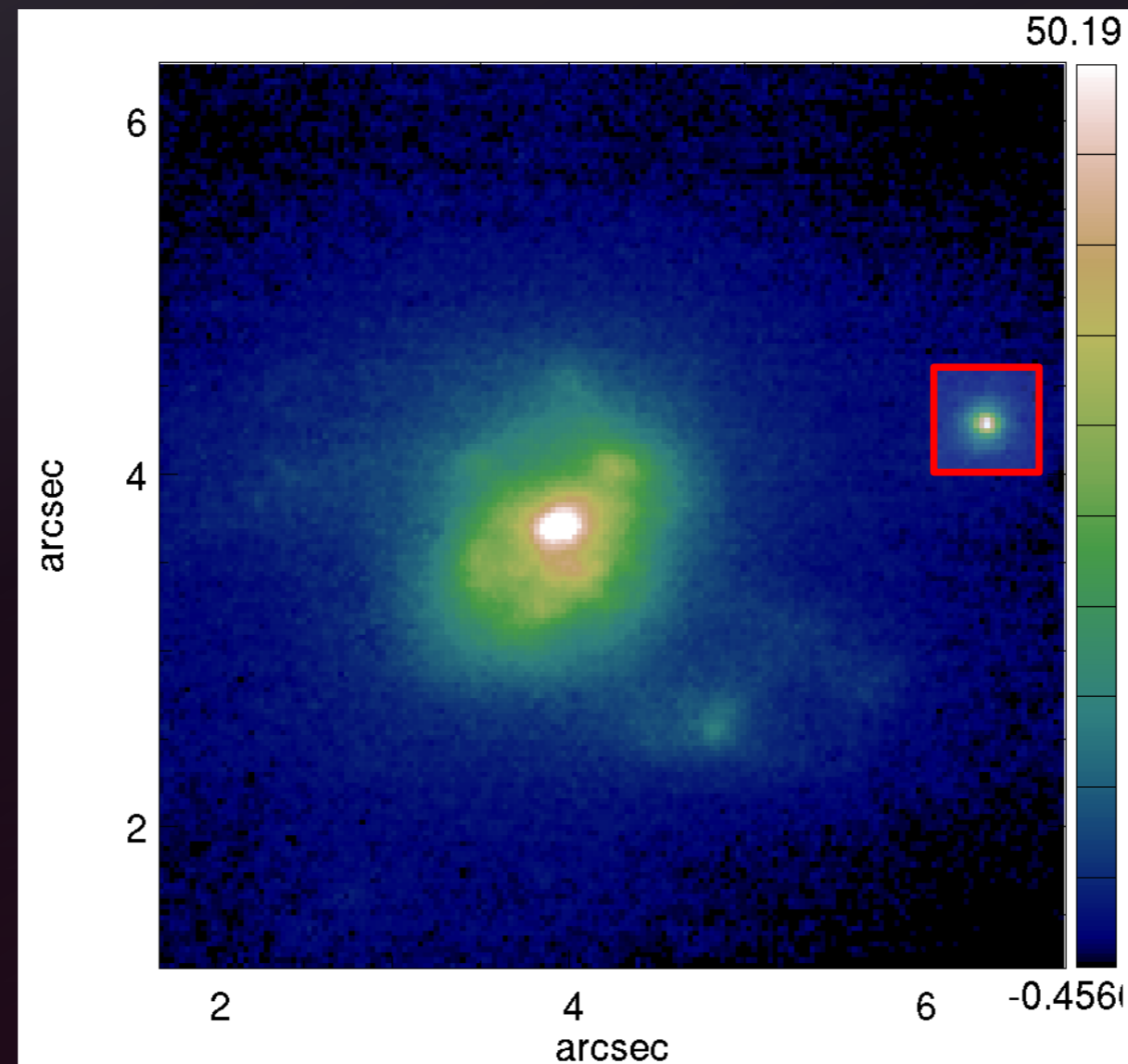
MOAO performance

- ◉ Comparing CANARY data to archive data
- ◉ IRAS21101+5810 : observed with HST



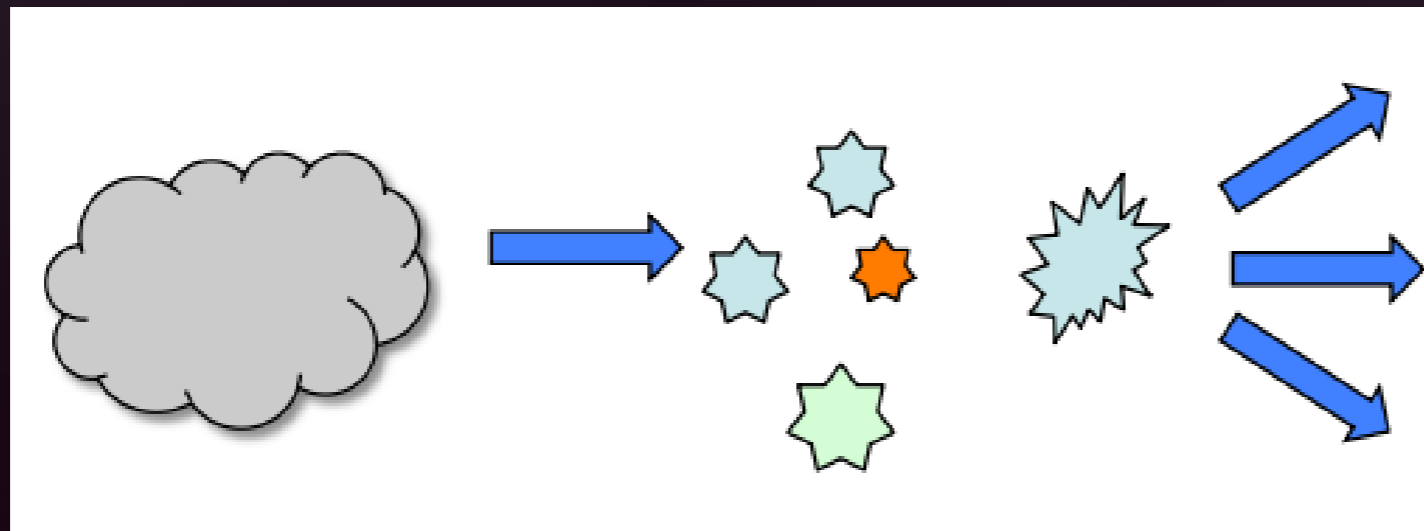
MOAO performance

- Field star on IRAS21101+5810 : monitor image quality during a long (>1h) exposure
 - FWHM varies from 0.12" to 0.18" (diffraction on WHT @ H : 0.08")



Preliminary science analysis

- Super Star Clusters : extreme star formation
 - Plays a major rôle in galaxy structural and chemical evolution
 - Massive conversion of gas into stars, heavy elements production
 - Distribution in the ISM through post-AGB and SNe
 - Power injection in the galactic medium and interaction with central black hole (feeding, feed-back)



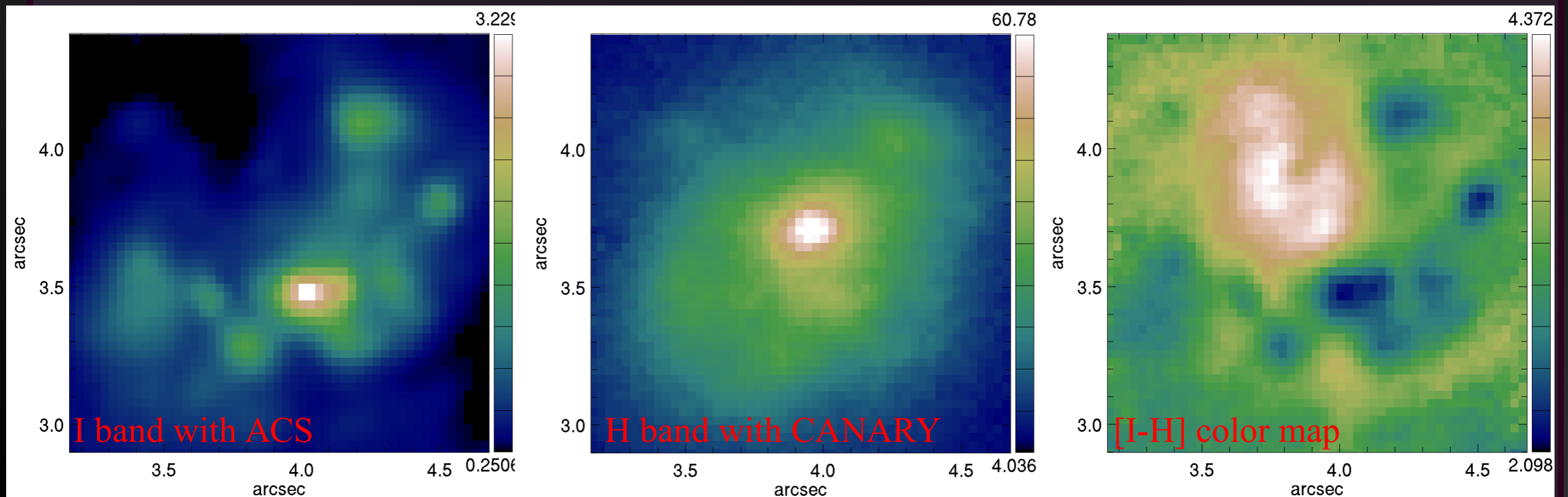
Halo, disk, bulge

Power

Heavy elements

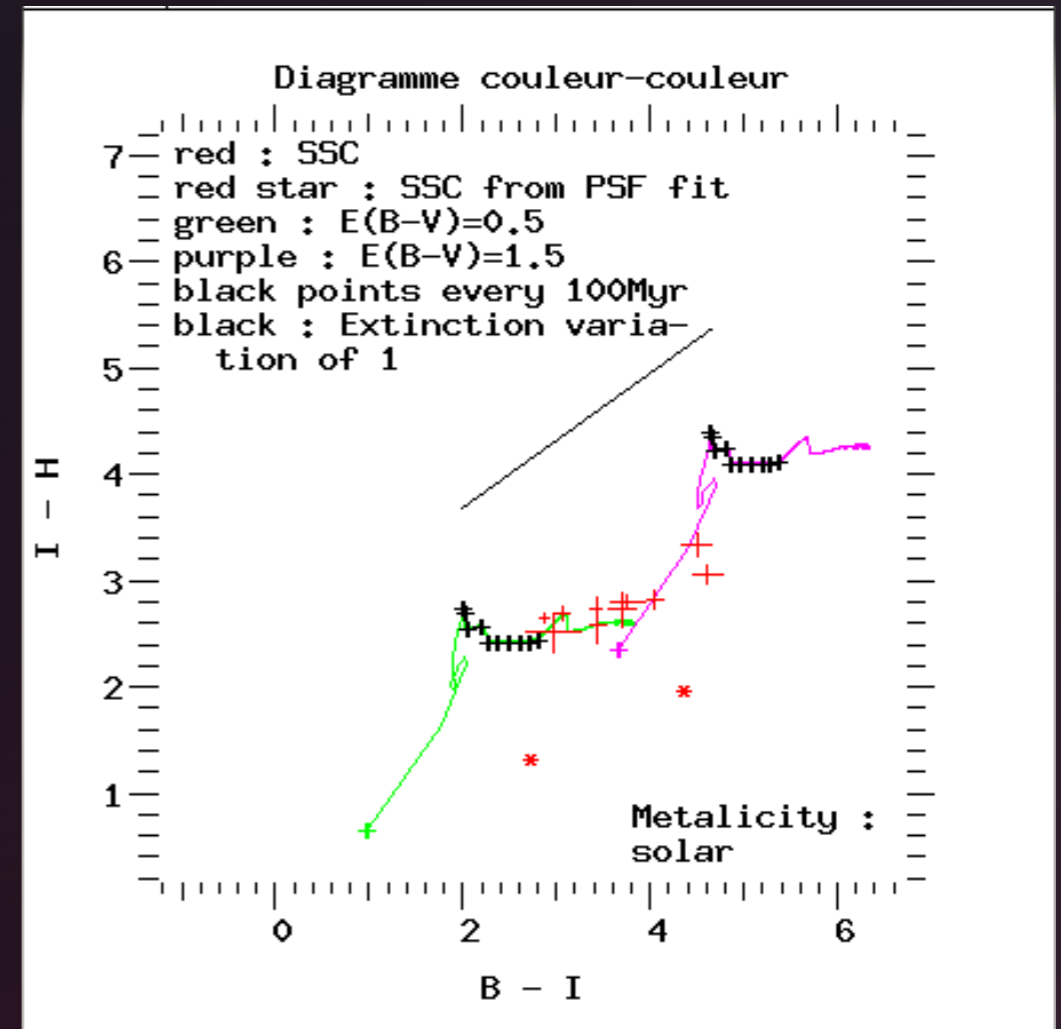
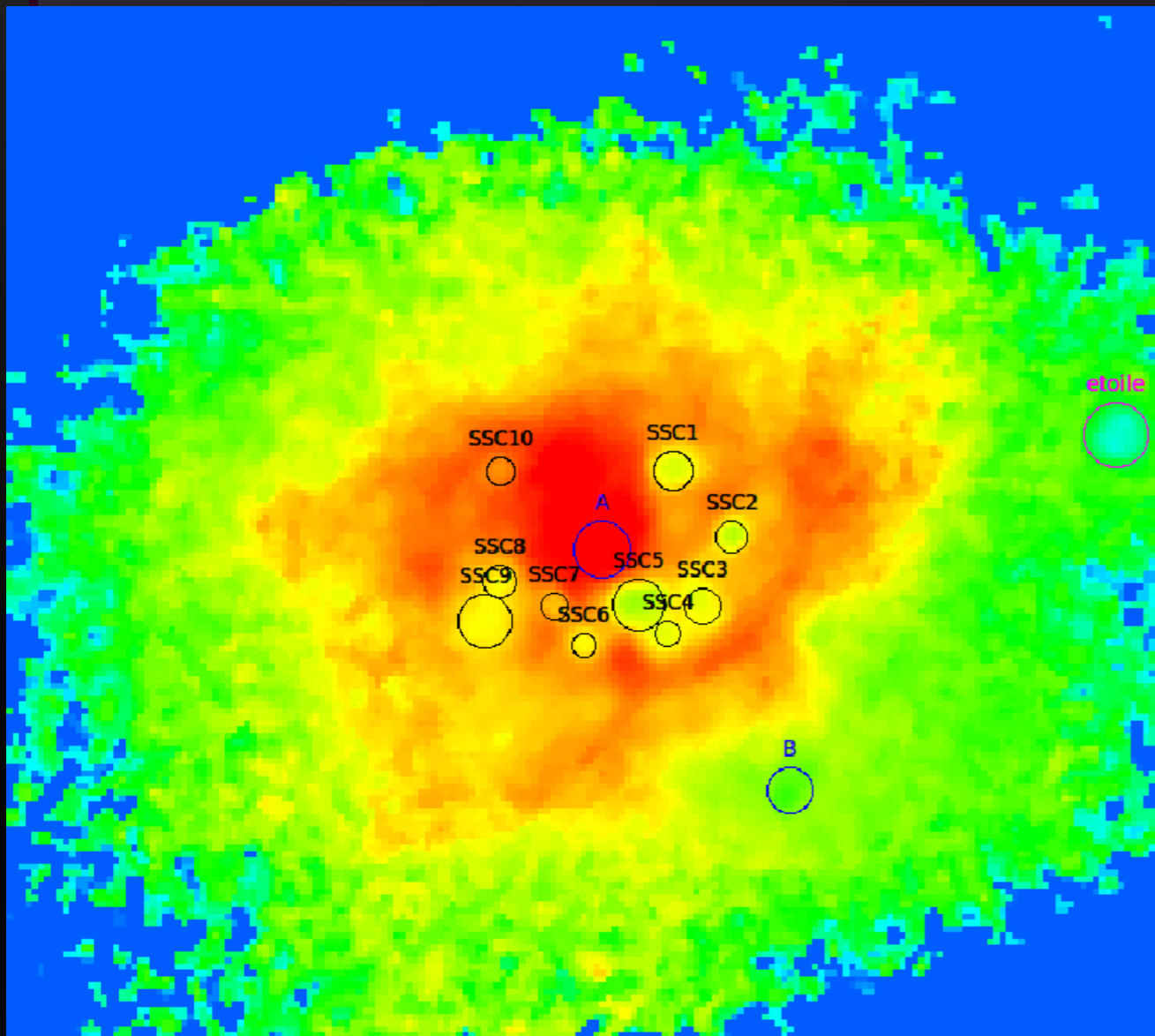
Preliminary science analysis

- ◉ Multi-spectral analysis : measure the SED of identified SSCs and derive age and physical properties
- ◉ The image quality obtained in the near-IR with CANARY allows a direct comparison with HST visible



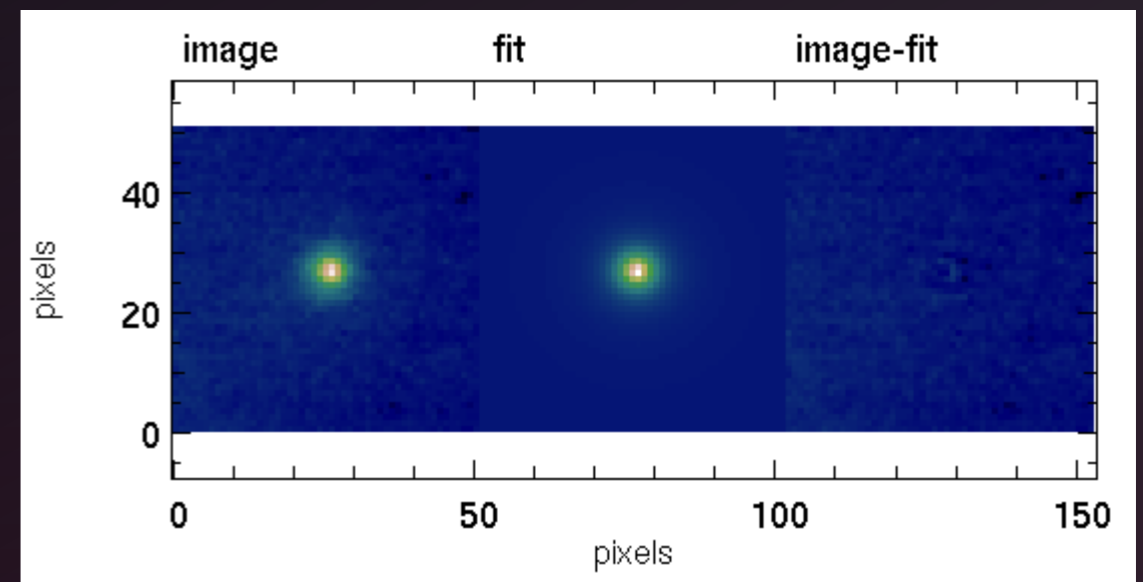
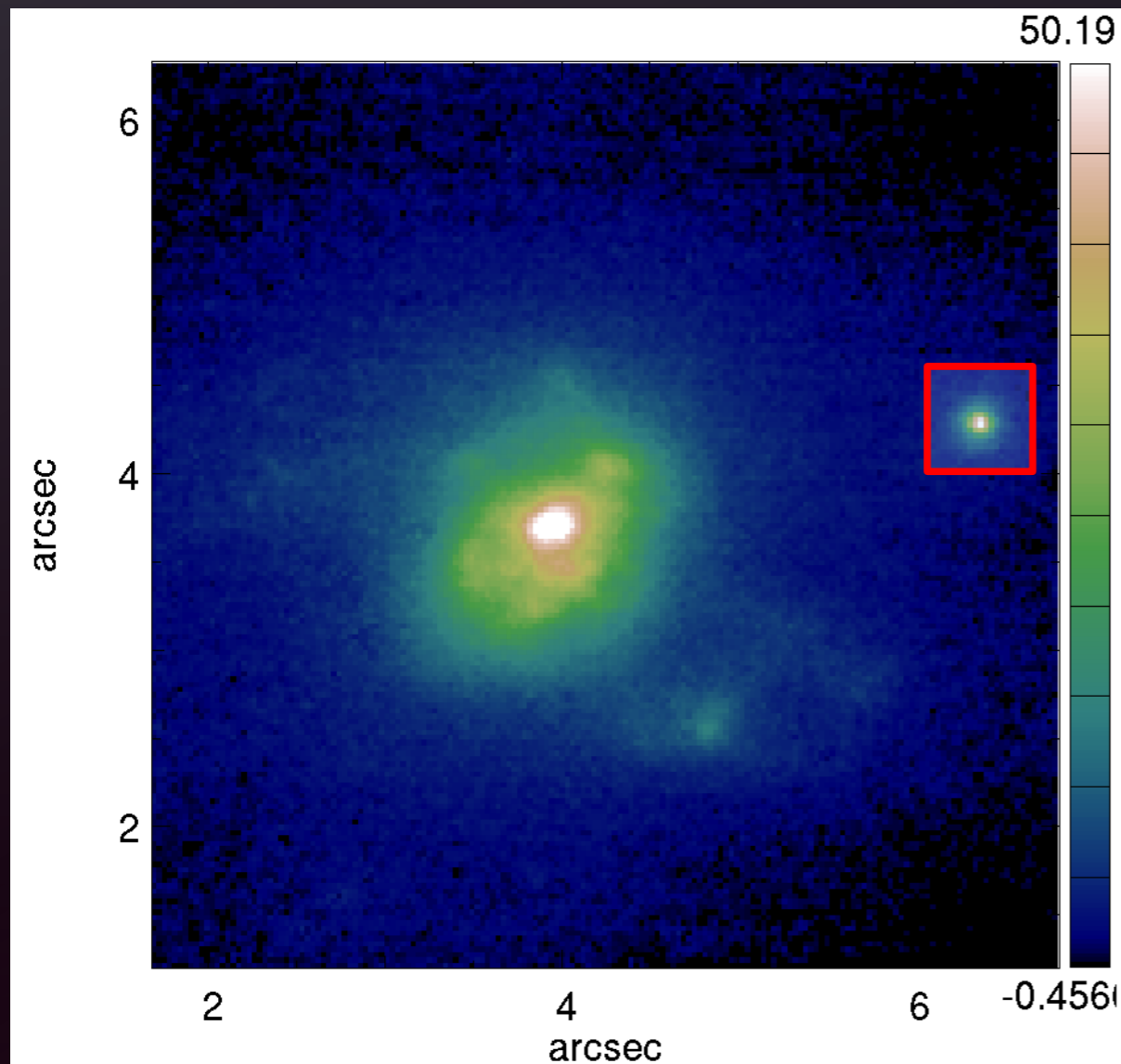
Preliminary science analysis

- ◉ Multi-spectral analysis : measure the SED of identified SSCs and derive age and physical properties
- ◉ Strong background in the IR : photometry with PSF fitting very difficult



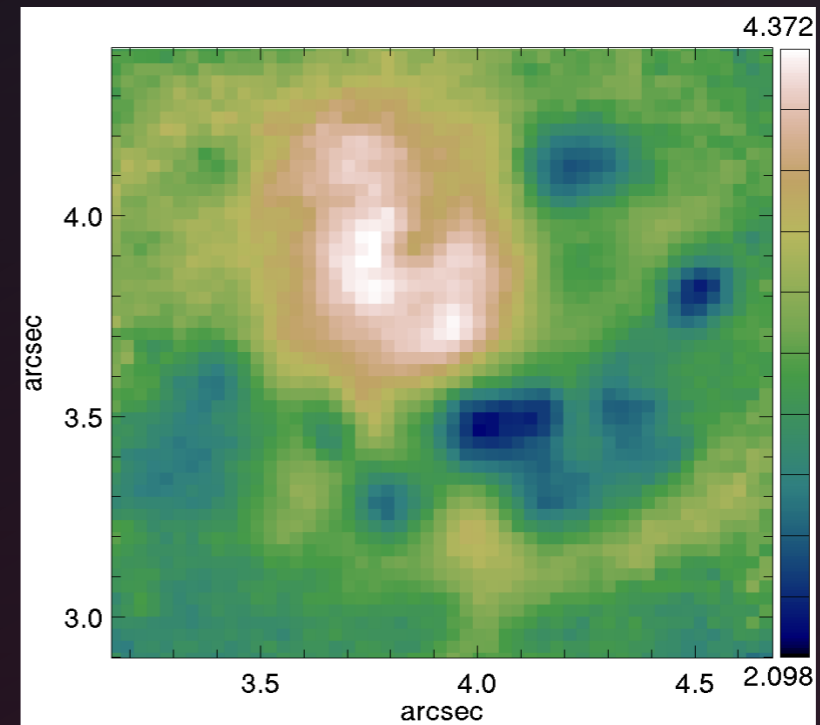
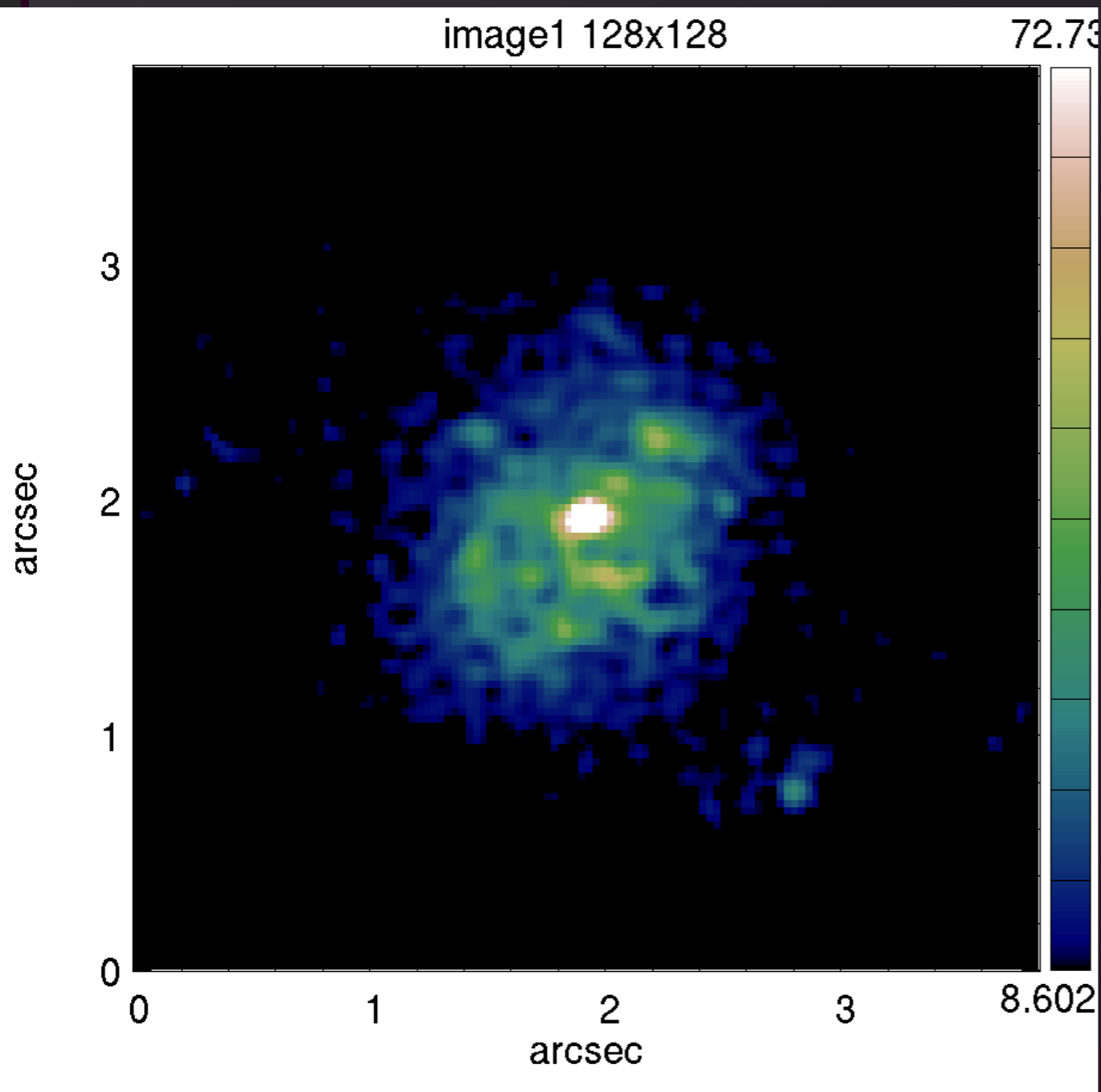
Preliminary science analysis

- ◉ Advanced data processing : deconvolution ?
- ◉ Using a model of the PSF measured on the image (Spydr Yorick plugin)



Preliminary science analysis

- Yoda deconvolution algorithm (maximum likelihood approach « à la MISTRAL » : <https://github.com/dgratadour/Yoda>)



Conclusions

- ◉ MOAO provides adequate performance for science
 - ◉ Image quality in the near-IR on a 4m ground based telescope comparable to HST visible
 - ◉ Stable performance over long exposures
- ◉ CANARY has limited science capabilities
 - ◉ Near-IR camera with a set of broad and narrow band filters
 - ◉ Not designed for science observations (pointing, interaction with telescope, ...)
- ◉ RAVEN has spectroscopic capabilities on a 8m telescope
 - ◉ Very powerful for physical diagnostics
- ◉ CANARY science team welcomes good ideas for new programs or collaborations !

Can't wait for MOAO on the E-ELT !

